

AN IMAGING OBSERVATION OF SN 1987A AT GAMMA-RAY ENERGIES

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ABSTRACT

The Caltech imaging γ -ray telescope was launched by balloon from Alice Springs, NT, Australia for observations of SN 1987A during the period 1987 November 18.60–18.87 UT. The preliminary results presented here are derived from 8200 s of instrument live time on the supernova and 2500 s on the Crab Nebula and pulsar at a float altitude of 36 km. We have obtained the first images of the SN 1987A region at γ -ray energies confirming that the bulk of the γ -ray emission comes from the supernova and not from LMC X-1. We compare our flux measurements to recent predictions concerning the distribution of ^{56}Co in the supernova ejecta and find the data to be consistent with models invoking moderate mixing of core material into the envelope.

Subject headings: gamma rays: general — gamma rays: spectra — stars: individual (SN 1987A) — stars: supernovae

I. INTRODUCTION

The occurrence of SN 1987A in the Large Magellanic Cloud has provided the first opportunity to study the hard X-ray and γ -ray emission from a Type II supernova. There have been numerous predictions of the γ -ray line flux from Type II events (cf. Bodansky, Clayton, and Fowler 1968; Clayton, Colgate, and Fishman 1969; and Colgate and McKee 1969); in contrast, the continuum flux has received less attention until the recent SN 1987A event. Initial hard X-ray and γ -ray detections of SN 1987A were made by several satellite instruments during 1987 July–August (Sunyaev *et al.* 1987, 1988; Dotani *et al.* 1987; and Matz *et al.* 1988). Further observations were performed by several balloon-borne instruments during the period 1987 October–December. This *Letter* reports on one of these—the first imaging observation of the supernova at γ -ray energies.

The observations were performed with the Caltech Gamma-Ray Imaging Payload (GRIP), a balloon-borne coded-aperture telescope sensitive to radiation in the energy range from 30 keV to 10 MeV (Althouse *et al.* 1985). The instrument employs a rotating hexagonal-celled uniformly redundant array (URA) and a 5 cm thick by 41 cm diameter position sensitive NaI scintillation detector to image a 14° diameter field of view with $1^\circ 1'$ angular resolution. The mask cell size used in the flight on 1987 November 18 was 4.8 cm, and the separation of the mask and detector was 2.5 m. The instrument has flown twice previously (Althouse *et al.* 1987), including a flight in 1987 May from Australia to observe SN 1987A. No γ -ray or hard X-ray emission was detected in the 1987 May observation (Cook *et al.* 1987).

II. IMAGES

The analysis of data from a coded-aperture detector involves the detection of a source and determination of its position, followed by detailed analysis of the spectral characteristics of the excess flux at the position of the source. We will briefly summarize the analysis procedure, which is described in more detail by Cook *et al.* (1984).

Photon event positions are obtained using a position-determination algorithm derived from laboratory calibrations.

For a given orientation of the coded-aperture mask, each event is back-projected onto the sky through the center of each mask cell to create a sky image, contributing a positive count to the sky image if a cell is open, a negative count if the cell is obscured with a lead block, and, for technical reasons, zero counts if the cell is the center cell of the hexagonal URA pattern. Because the number of open and blocked cells is equal, the sum of counts in the sky image is zero.

The data from different mask orientations (typically 360 in our analysis) are then added by rotating the individual sky images to a common orientation; this addition also implements the mask-antimask background subtraction for the hexagonal URAs. Finally, the resulting image is smoothed by convolution with an angular point-spread function derived from the detector position-determination point-spread function and the size and shape of the mask cells. This yields an image which is statistically optimal for measurements of the flux and location of point sources (Cook *et al.* 1984).

The convolved image is calibrated in terms of the statistical significance of the excess number of counts at a given image position. The statistics of the image are the Poisson statistics of the count difference between events passing through open and blocked mask cells and are dependent only on the total number of events and the angular point-spread function employed in the convolution.

Given a sky image over a broad energy range, the count distribution is scanned for significant excesses. If a source is detected, the source spectrum is determined from images for several different energy bands. The flux measurement for each energy band is derived from the count excess (positive or negative) at the position of the source, taking into account the efficiency factor appropriate for the convolving function, the observation live time, the detector's effective area, and the attenuation of photons in the atmosphere and instrument.

Figure 1 shows an image of the SN 1987A region for energies between 40 and 1300 keV. The image is a composite of data from four individual supernova observation periods, which alternated with four calibration observations of the Crab Nebula and pulsar. The supernova image fields were aligned by using residual pointing corrections determined from the

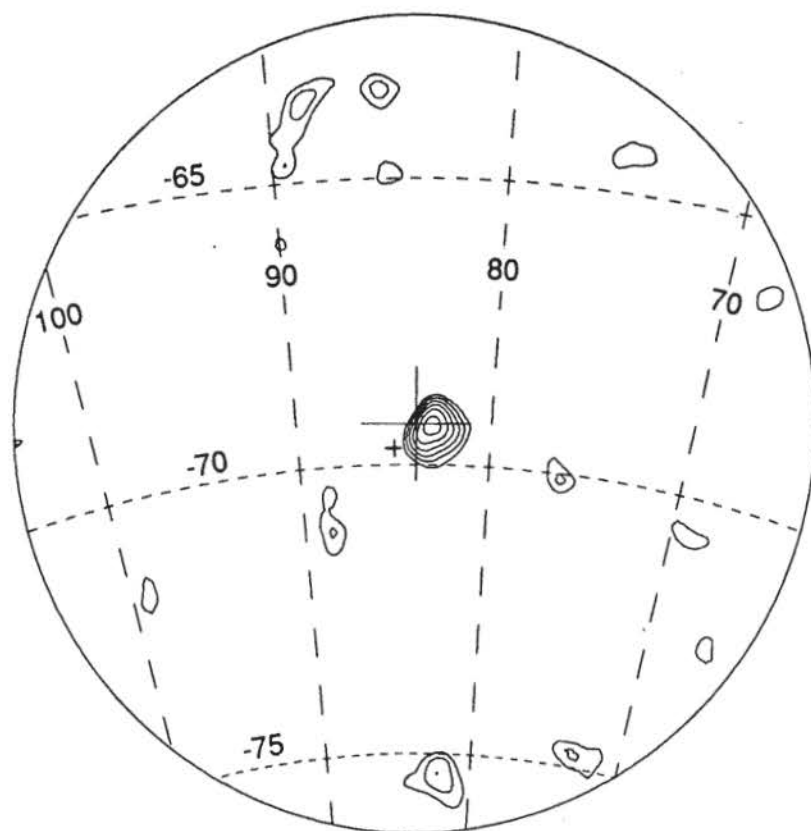


FIG. 1.—Image of the SN 1987A region from 40 to 1300 keV covering a 14° field of view. Right ascension (vertical lines) and declination (horizontal lines) are indicated. The contours indicate the number of excess counts in a given direction calibrated in units of the statistical significance of the excess, with contours beginning at the 2σ level and spaced by 0.5σ intervals. The large cross indicates the expected position of SN 1987A. The small cross indicates the expected position of LMC X-1.

Crab observations. Flux is detected at a nominal 4.9σ level from within 0.28° of the expected position of SN 1987A. The 1σ error in the peak location is estimated to be 0.14° in the elevation direction and 0.27° in the azimuthal direction. The errors include both statistical error and an estimate of residual systematic errors determined from the Crab observations. The deviation from the expected position is in the azimuthal direction, and the larger error in this direction may be due to deviations between the actual and the assumed magnetic field direction used as a pointing reference during flight. In the elevation direction, which is approximately aligned with the declination axis in these observations, the peak centroid is consistent with the position of the supernova.

Because these are the first imaging observations of SN 1987A at energies greater than 30 keV, it is valuable to investigate other astrophysical sources as the possible origin of the gamma-ray emission. The one nearby object most likely to be a contaminating source is the black hole candidate LMC X-1 at $\alpha(1950) = 85^\circ 03'$ and $\delta(1950) = -69^\circ 77'$, which is 0.60° from SN 1987A at $\alpha(1950) = 83^\circ 96'$ and $\delta(1950) = -69^\circ 30'$. It is conceivable that LMC X-1 could have a hard, time-variable, non-thermal tail of emission extending up to MeV energies, much like Cyg X-1. In fact, LMC X-1 is known to be a strong source at X-ray energies, and there is some uncertainty over its possible contribution to the X-ray flux reported from SN 1987A (Sunyaev *et al.* 1987; Dotani *et al.* 1987).

Because of the azimuthal uncertainty in our aspect determination at the current stage of analysis, it is impossible to eliminate LMC X-1 entirely as a contributor to the gamma-ray flux.

However, because LMC X-1 is 0.5° away from SN 1987A in the elevation direction, and because the 1σ error for the position determination in the elevation direction is 0.14° , the chance that LMC X-1 is the sole source of emission is less than 0.1%. For the remainder of this Letter, we assume that the supernova is the sole source of excess photons in the peak in the image of Figure 1.

III. SPECTRA

Measurements of the γ -ray flux from SN 1987A for selected energy intervals are listed in Table 1 and shown in Figure 2. The flux measurements were derived from the number of excess counts in the specified energy interval at the peak position

TABLE 1
SN 1987A FLUX MEASUREMENTS

Energy Interval (keV)	Differential Flux ($10^{-5} \text{ cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$)	Integral Flux ($10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$)
51–101	4.73 ± 2.17	...
101–173	2.29 ± 0.84	...
173–303	1.04 ± 0.57	...
303–545	0.32 ± 0.47	...
545–801	0.33 ± 0.37	...
801–883	1.32 ± 0.69	$1.08 \pm 0.56^*$
883–1185	0.62 ± 0.38	...
1185–1281	1.15 ± 0.66	$1.11 \pm 0.64^*$
1281–1614	0.08 ± 0.43	...
1614–3752	$< 0.37 (3\sigma)$...

* Includes both continuum and line flux; see text for line flux estimates.

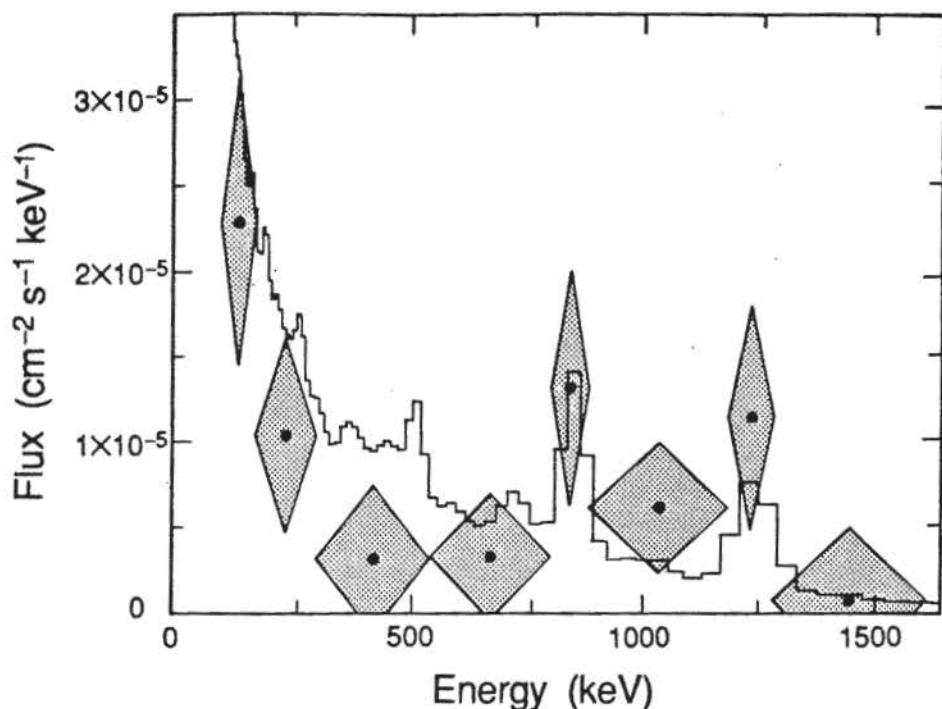


FIG. 2.—Flux measurements compared with the result of a Monte Carlo calculation of the expected flux on day 250 of the supernova (Pinto and Woosley 1988). The measurements are indicated by diamonds with the width indicating the energy interval of the measurement. The Monte Carlo histogram has been convolved with a Gaussian representation of the instrumental energy resolution.

determined from Figure 1. To obtain the overall characteristics of the spectrum, the data were fitted to a power-law differential continuum flux, $F = K(E/100 \text{ keV})^\alpha$, over the energy range 51–1614 keV. A very hard power-law index, $\alpha = -0.9$, and a flux normalization, $K = 2.8 \times 10^{-5} (\text{cm}^2 \text{ s keV})^{-1}$, were obtained. For comparison, the same type of analysis for the Crab yielded a flux of $4.4 \times 10^{-4} (E/100 \text{ keV})^{-2} (\text{cm}^2 \text{ s keV})^{-1}$ between 40 and 500 keV.

It is of interest to estimate the flux at the energies of the prominent γ -ray lines expected from ^{56}Co decay. Table 1 gives flux values for intervals including 847 keV and 1238 keV. The widths of the intervals are chosen such that 89% and 84%, respectively, of the narrow line flux is contained in the intervals.

To estimate the amount of possible ^{56}Co line flux enhancement above the continuum level, we have performed a minimum- χ^2 fit to our measurements in the energy range 545–1614 keV. For the fitting function we chose a straight-line continuum with gamma-ray line contributions described by two Gaussian functions centered at 847 and 1238 keV, having widths fixed at the instrumental energy resolution of 7% and 6.5% FWHM, respectively. The flux level and slope of the continuum contribution, as well as the line flux levels, were varied to minimize χ^2 . The resulting minimum occurred at a reduced χ^2 of 0.92 (with 26 degrees of freedom) for values of the 847 and 1238 keV line fluxes of $0.5 \pm 0.6 \times 10^{-3}$ and $1.3 \pm 0.8 \times 10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$, respectively.

The results reported here are preliminary. Most conceivable systematic errors, in particular, pointing uncertainties, would tend to yield an absolute measured flux somewhat lower than the true flux. We note that the estimate of the Crab spectrum given above is $\sim 20\%$ – 40% below other measurements at 100 keV (Jung 1986) and may indicate an overall systematic underestimate of the flux. However, the earlier measurements

were made with nonimaging instruments which may overestimate the flux, and there is also the possibility of time variability.

IV. DISCUSSION

The results from Table 1 should be compared to other measurements of the hard X-ray and γ -ray flux made during the period of 1987 August through November. These include results using the *Mir-Kvant* observatory (Sunyaev *et al.* 1987, 1988), the *Ginga* satellite (Dotani *et al.* 1987), and the *Solar Maximum Mission* (SMM) observatory (Matz *et al.* 1988). Our continuum flux measurements on 1987 November 18 at energies from 51 to 173 keV are consistent with the flux of approximately $3.0 \times 10^{-5} (\text{cm}^2 \text{ s keV})^{-1}$ at 100 keV measured by the *Mir-Kvant* observatory during 1987 November (Sunyaev *et al.* 1988). The integral fluxes for the 847 keV and 1238 keV regions from Table 1 include both line and continuum contributions. The fits to the lines discussed above yielded values of $0.5 \pm 0.6 \times 10^{-3}$ and $1.3 \pm 0.8 \times 10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$ for 847 and 1238 keV, respectively, consistent with the γ -ray line fluxes reported by the SMM instrument averaged over the 1987 August through October period: $1.0 \pm 0.25 \times 10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$ at 847 keV and $0.6 \pm 0.2 \times 10^{-3} \text{ cm}^{-2} \text{ s}^{-1}$ at 1238 keV. Our measurements should also be compared to the forthcoming results from three other balloon-borne instruments. These include the hard X-ray measurements from 45 to 200 keV of Wilson *et al.* (1988) on 1987 October 29–30, and the high-resolution γ -ray line measurements of Sandie *et al.* (1988) and Mahoney *et al.* (1988) on 1987 October 29–30 and 1987 December 6, respectively.

A likely origin of the observed hard X-ray and γ -ray flux is the decay of ^{56}Co which is a decay product of ^{56}Ni synthesized in the supernova explosion (Bodansky, Clayton, and Fowler 1968; Clayton, Colgate, and Fishman 1969; Colgate and

McKee 1969). The ^{56}Co decays are expected to yield both line emission and a continuum of Compton-degraded photons. Approximately $0.075 M_{\odot}$ of ^{56}Ni was synthesized in the initial explosion, as determined from the decay of the bolometric luminosity of the supernova (Woosley 1988). Results from Table 1 are compared with recent theoretical predictions in Figure 2. The continuous histogram is based on a Monte Carlo calculation of the expected photon flux on day 250 (18 days before our observation), taken from Pinto and Woosley (1988). Similar calculations have been carried out by other authors (Itoh *et al.* 1987; McCray, Shull, and Sutherland 1987; Arnett and Fu 1989; Ebisuzaki and Shibazaki 1988; Fu and Arnett 1989; Shibazaki and Ebisuzaki 1988; Xu *et al.* 1988); the calculation in Figure 2 is an example of models which incorporate mixing to explain the early turn-on of X-rays and γ -rays seen by the *Mir-Kvant*, *Ginga*, and *SMM* instruments. The model assumes a $6 M_{\odot}$ He core, a $10 M_{\odot}$ hydrogen envelope,

1.3×10^{51} ergs explosion kinetic energy, and a moderate degree of mixing of ^{56}Co out into the helium envelope. Our flux measurements shown in Figure 2 are in qualitative agreement with the calculations. In particular, when the flux in the regions of the lines is compared to the continuum flux measurements, the results suggest a flux excess in the region of the expected γ -ray lines from ^{56}Co decay as predicted by the Monte Carlo calculations for this time period.

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